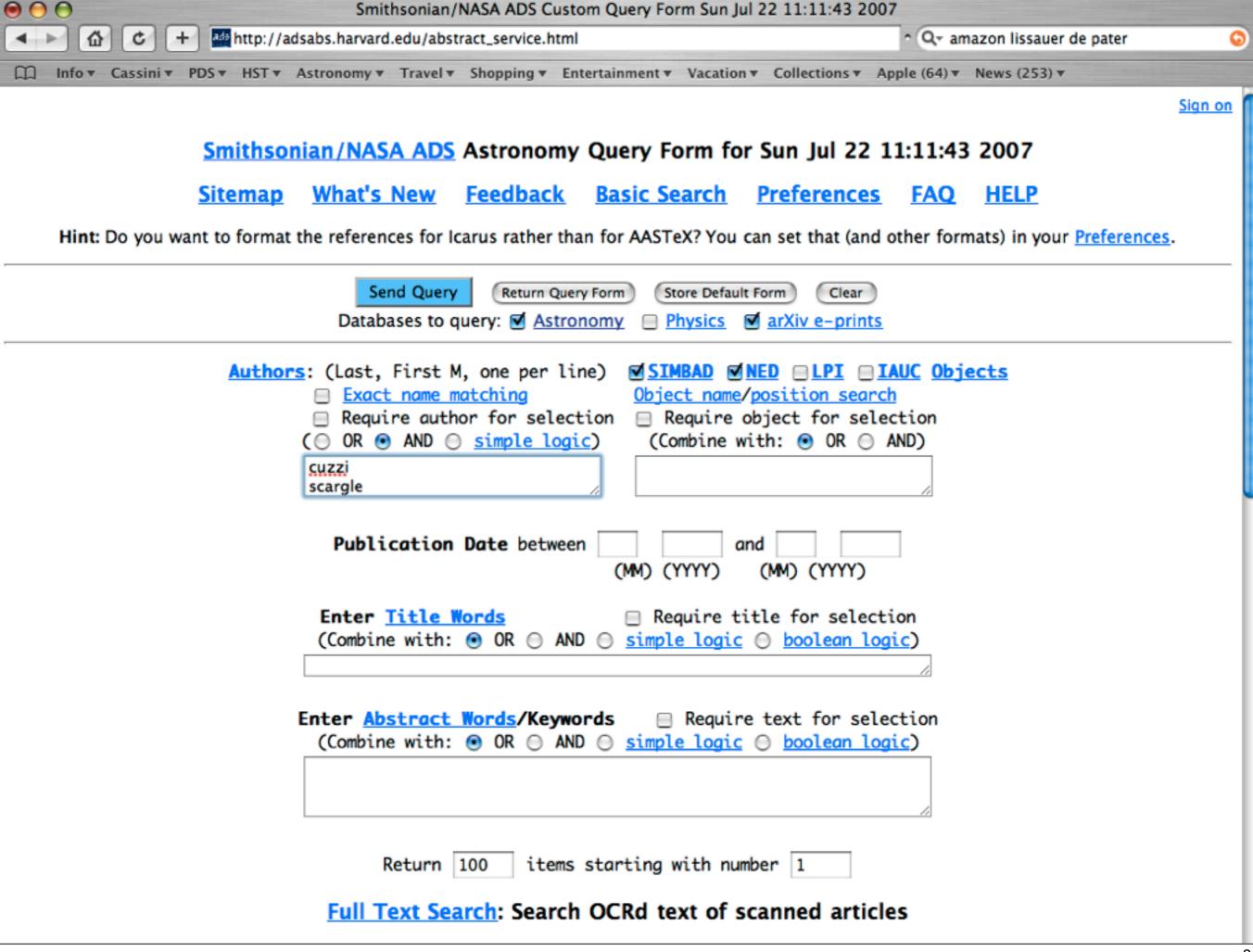
Planetary Ring Studies: Tools of the Trade

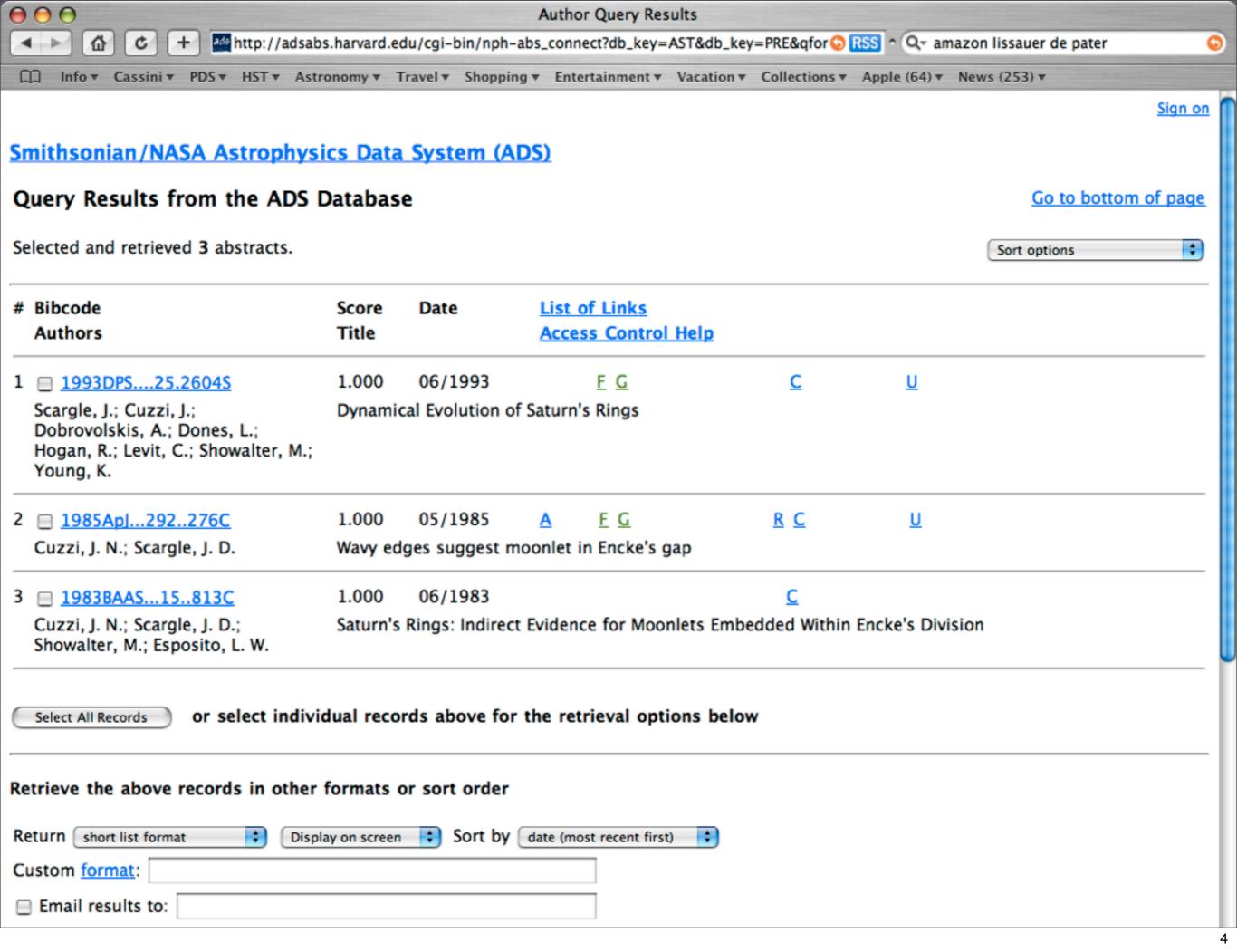
Mark R. Showalter SETI Institute

COSPAR WORKSHOP Wednesday, July 23, 2007

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Translate Abstract

Title: Wavy edges suggest moonlet in Encke's gap

Authors: Cuzzi, J. N.; Scargle, J. D.

Affiliation: AA(NASA, Ames Research Center, Space Science Div., Moffett Field, CA), AB(NASA, Ames Research Center, Space Science

Div., Moffett Field, CA)

Publication: Astrophysical Journal, Part 1 (ISSN 0004-637X), vol. 292, May 1, 1985, p. 276-290. (Apj Homepage)

Publication Date: 05/1985

Category: Lunar and Planetary Exploration; Saturn

Origin: STI: LPI [AN-850775%]]

NASA/STI Keywords: PLANETOLOGY, SATURN RINGS, SATURN SATELLITES, IMAGE ANALYSIS, POWER SPECTRA, VOYAGER 1 SPACECRAFT,

VOYAGER 2 SPACECRAFT, WAVELENGTHS

LPI Keywords: SATURN, ENCKE DIVISION, MOONLETS, VOYAGER MISSIONS, IMAGERY, FEATURES, A RING, WAVES, LONGITUDES, POSITION

(LOCATION), WAVELENGTHS, SIZE, RADIUS, ORBITS, OBSERVATIONS, SHEPHER SATELLITES, SATELLITES

DOI: <u>10.1086/163158</u>

Bibliographic Code: 1985ApJ...292..276C

Abstract

Voyager images have revealed radial undulations of the inner and outer edges of the 325 km wide Encke gap in Saturn's A ring. These waves are present at some, but not all, longitudes. Their locations and wavelengths provide strong indirect evidence for the presence of at least one dominant moonlet of about 10 km radius orbiting near the center of the gap. Implications for 'shepherding' theory are discussed.

Printing Options

5

THE ASTROPHYSICAL JOURNAL, 292: 276-290, 1985 May 1 © 1985. The American Astronomical Society. All rights reserved. Printed in U.S.A.

WAVY EDGES SUGGEST MOONLET IN ENCKE'S GAP

JEFFREY N. CUZZI AND JEFFREY D. SCARGLE Space Science Division, NASA/Ames Research Center, Moffett Field, California Received 1984 September 4; accepted 1984 November 14

ABSTRACT

Voyager images have revealed radial undulations of the inner and outer edges of the 325 km wide Encke gap in Saturn's A ring. These waves are present at some, but not all, longitudes. Their locations and wavelengths provide strong indirect evidence for the presence of at least one dominant moonlet of about 10 km radius orbiting near the center of the gap. Implications for "shepherding" theory are discussed.

Subject headings: planets: satellites — planets: Saturn

I. INTRODUCTION

The complex structure discovered in the rings of Saturn by the Voyager spacecraft has provided a fertile testing ground for a variety of dynamical theories. Especially within the outer (A) ring, gravitational resonances with known satellites external to the rings are of sufficient strength and abundance to account for most of the observed structure (Lissauer and Cuzzi 1982; Holberg, Forrester, and Lissauer 1982; Esposito et al. 1984). However, resonances with external satellites are incapable of accounting either for the "record-groove" appearance of the B ring or for the existence of a handful of essentially empty gaps with widths between 50 and 400 km which occur in the A, C, and Cassini Division regions of the rings (Cuzzi et al. 1984).

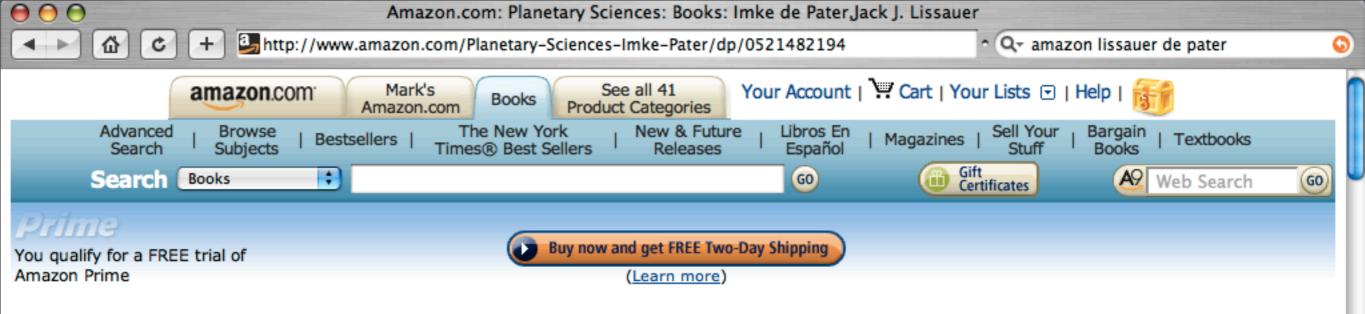
The existence of a population of embedded "moonlets" has been invoked to explain these latter effects. A moonlet exerts a torque on adjacent ring material which results in transfer of momentum to or from the ring (Lin and Papaloizou 1979,

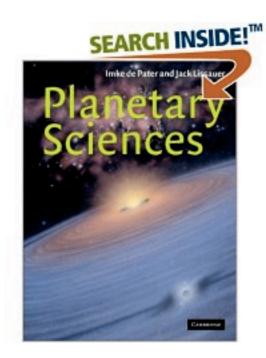
The physics of this kinking was developed by Julian and Toomre (1966) and described also by Lin and Papaloizou (1979) and Dermott (1981). Encke's gap contains several narrow ringlets qualitatively similar to the F ring; in this paper we report the observation and analysis of edge waves running along the inner and outer edges of the Encke gap. These waves are more easily understood than the kinky ringlets, and strongly support the hypothesis that the Encke gap contains at least one, as yet unseen, embedded moonlet of roughly 10 km radius. In subsequent papers we will present complementary evidence and studies of the kinky ringlets in the Encke gap. We will also systematically explore the edges of other empty gaps. Preliminary, very cursory, investigation has revealed no wavy edges in other gaps such as Huygens or Maxwell.

Section II deals with the observations (Voyager images) and techniques for determining wave properties. Section III briefly reviews the relevant physics of local ring-moon interactions. In § IV we compare the observations with theoretical expecta-

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- Murray & Dermott: Solar System Dynamics
 - The best reference on all aspects of planetary dynamics.
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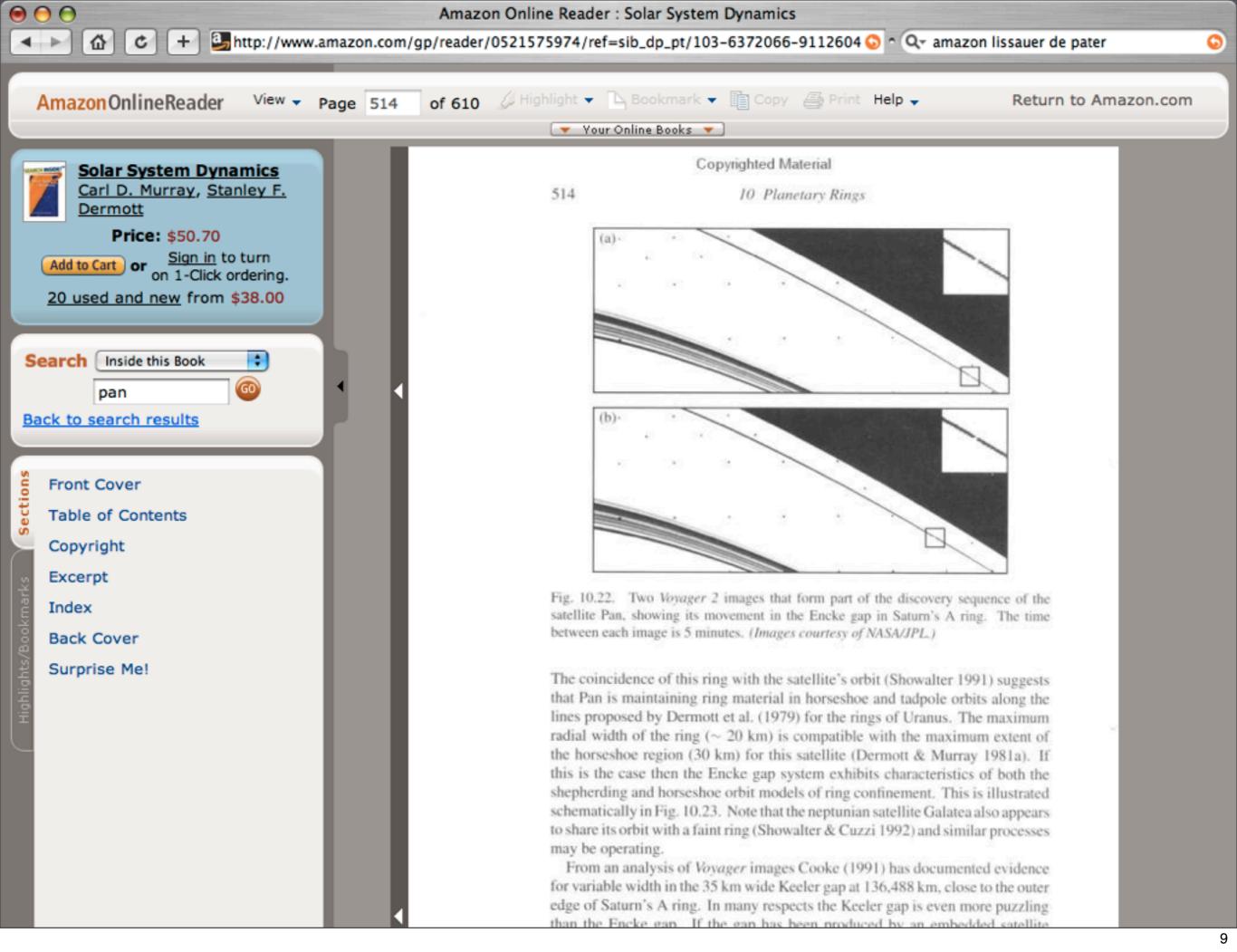


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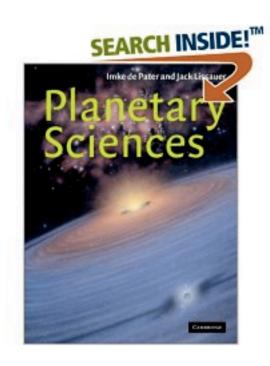
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SPICE Toolkit

- THE solution for planetary geometry associated with science products.
- FORTRAN, C or IDL.
- OUTSTANDING documentation.
- Not so bad! Really!
- Web site: naif.jpl.nasa.gov
- Data files: ftp naif.jpl.nasa.gov/pub/naif

fittp://naif.jpl.nasa.gov/naif/pds.html

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The Navigation and Ancillary Information Facility (NAIF) offers an information system named "SPICE" to assist scientists in planning and interpreting scientific observations from space-borne instruments. SPICE is also widely used in engineering tasks needing access to space geometry.

SPICE is focused on solar system geometry, time, and related information. The SPICE system includes a large suite of software, mostly in the form of subroutines, that customers use to read SPICE files and to compute derived observation geometry, such as altitude, lattitude/longitude, and lighting angles. SPICE data and software may be used within many popular computing environments. The software is offered in Fortran, C and IDL®, with a Matlab interface in the works.

SPICE is used on NASA's solar system exploration missions, and some NASA space physics and astrophysics missions. It is also being used as an adjunct to local national capabilities on some non-U.S. missions such as Mars Express, Rosetta, Venus Express and Hayabusa.

There is no charge to individuals to obtain SPICE data and software.

The export status with regard to SPICE components and services is provided under the RULES link on this website.

Announcements

 The IERS has announced there will *NOT* be a new leap second (positive or negative) declared for the December 2007 opportunity. Thus the current SPICE generic leapseconds kernel, named naif0008.tls, and any unmodified copy thereof, will remain valid through June 2008.





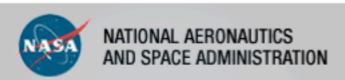




fttp://naif.jpl.nasa.gov/naif/toolkit_FORTRAN.html







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Home

About SPICE

About NAIF

Data

Toolkit

Documentation

Tutorials

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Support

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Feedback

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PC, Windows, Compaq Visual (Digital) FORTRAN

PC, Windows, Intel FORTRAN

PC, Windows, Lahey FORTRAN95

Sun, Solaris, SUN FORTRAN



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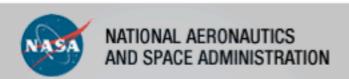


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Home

About SPICE

About NAIF

Data

Toolkit

Documentation

Tutorials

Lessons

Support

Rules

Feedback

Site Map

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PC, Linux, gCC

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PC, Windows, Microsoft Visual C

Sun, Solaris, SUN C

Sun, Solaris, gCC

Sun, Solaris, gCC/64bit



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Home

About SPICE

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Data

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Documentation

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Lessons

Support

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Feedback

Site Map

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Mac/Intel, OS-X, Apple C/IDL 6.3

PC, Linux, qCC/IDL 6.3

PC, Windows, Microsoft Visual C/IDL 6.3

Sun, Solaris, SUN C/IDL 6.3

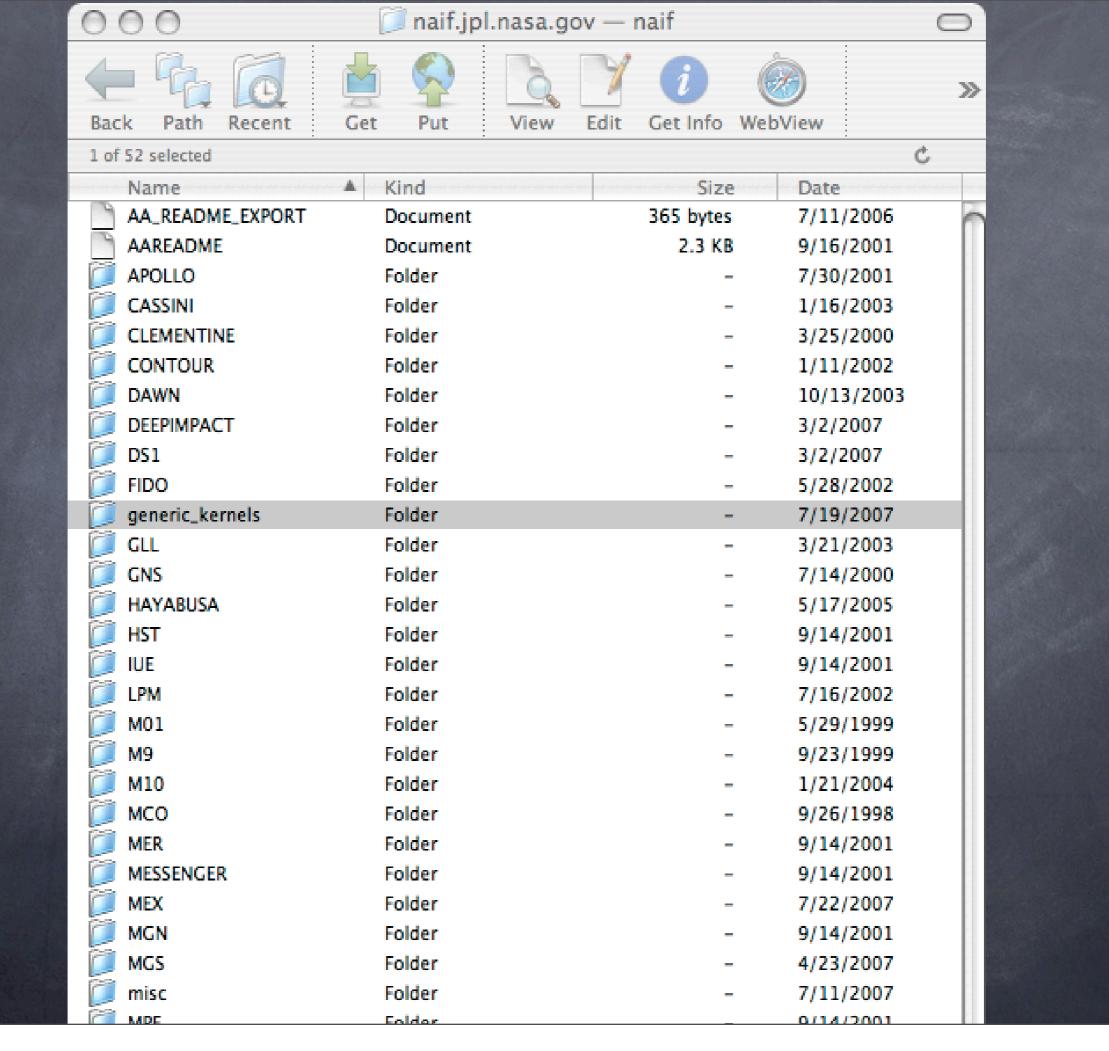
Sun, Solaris, qCC/IDL 6.3

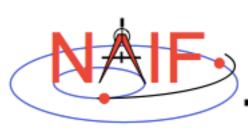


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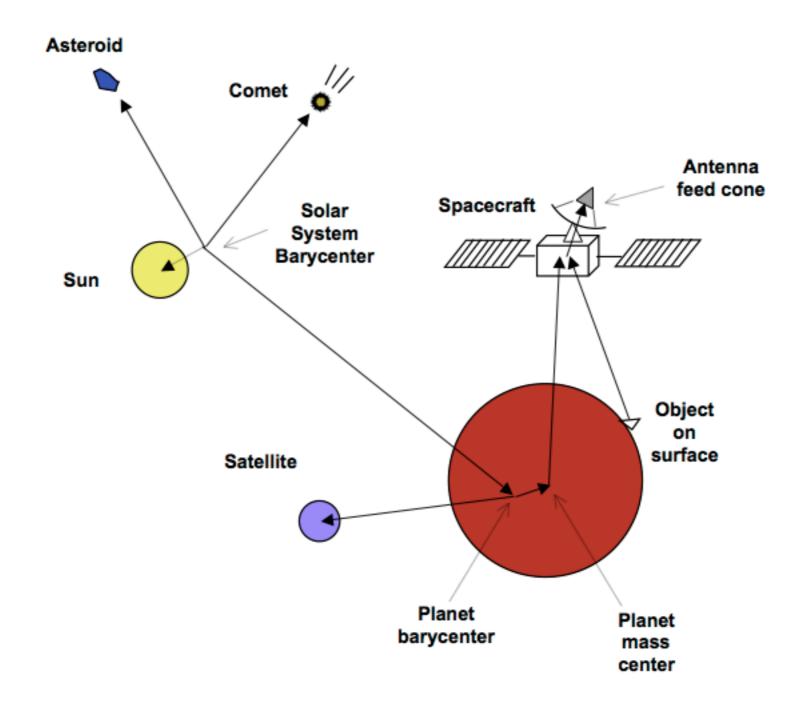
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Examples of SPICE Ephemeris Objects

Navigation and Ancillary Information Facility



SPK Subsystem





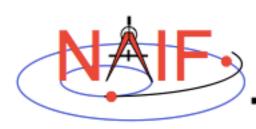












Reading an SPK File

Navigation and Ancillary Information Facility

Initialization...typically done once per program execution

Tell your program which SPICE files to use ("loading" files)

CALL FURNSH ('spk_file_name')

CALL FURNSH ('leapseconds_file_name')

Better yet, replace these two calls with a single call to a "furnsh kernel" containing the names of all kernel files to load.

Loop... do as many times as you need

Convert UTC time to ephemeris time (TDB), if needed CALL STR2ET ('utc_string', tdb)

Retrieve state vector from the SPK file at your requested time CALL SPKEZR (target, tdb, 'frame', 'correction', observer, state, lightime)

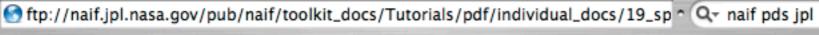


Use the returned state vector in computing geometry of interest.

SPK Subsystem











A Simple Example of Reading an SPK File

Navigation and Ancillary Information Facility

Initialization - typically do this just once per program execution

```
CALL FURNSH ( 'NAIF0008.TLS' )
                                                  Better to use a "furnsh kernel"
                                                  instead of these individual
CALL FURNSH ( 'HUYGENS_3_MERGE.BSP'
                                                  FURNSH statements
```

Repeat in a loop as needed to solve your particular problem

```
CALL STR2ET ('2004 NOV 21 02:40:21.3', TDB )
CALL SPKEZR ('TITAN', TDB, 'J2000', 'LT+S', 'HUYGENS PROBE',
              STATE, LT )
```

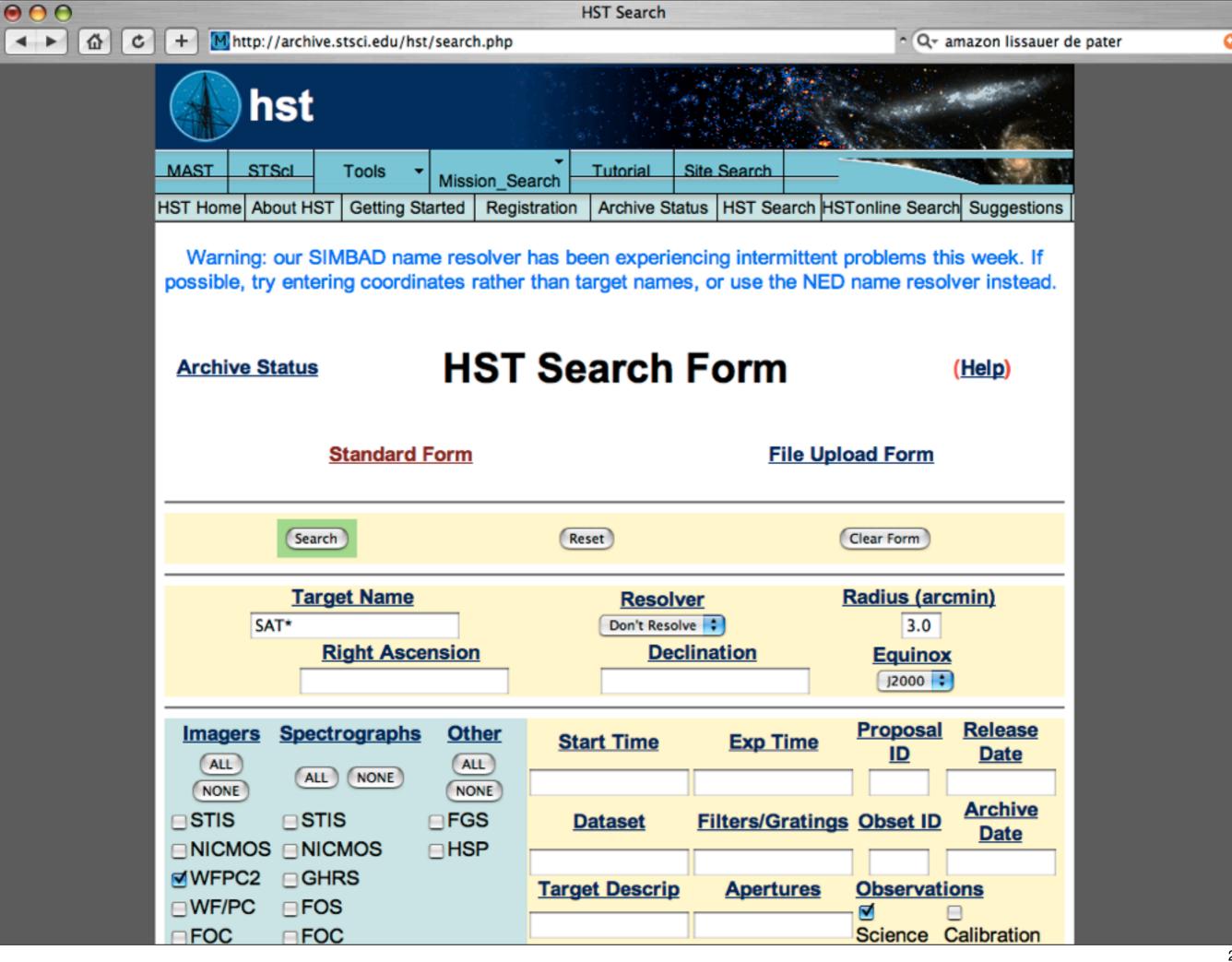
(Insert additional code here to make derived computations such as spacecraft sub-latitude and longitude, lighting angles, etc. Use more SPICE subroutines to help.)

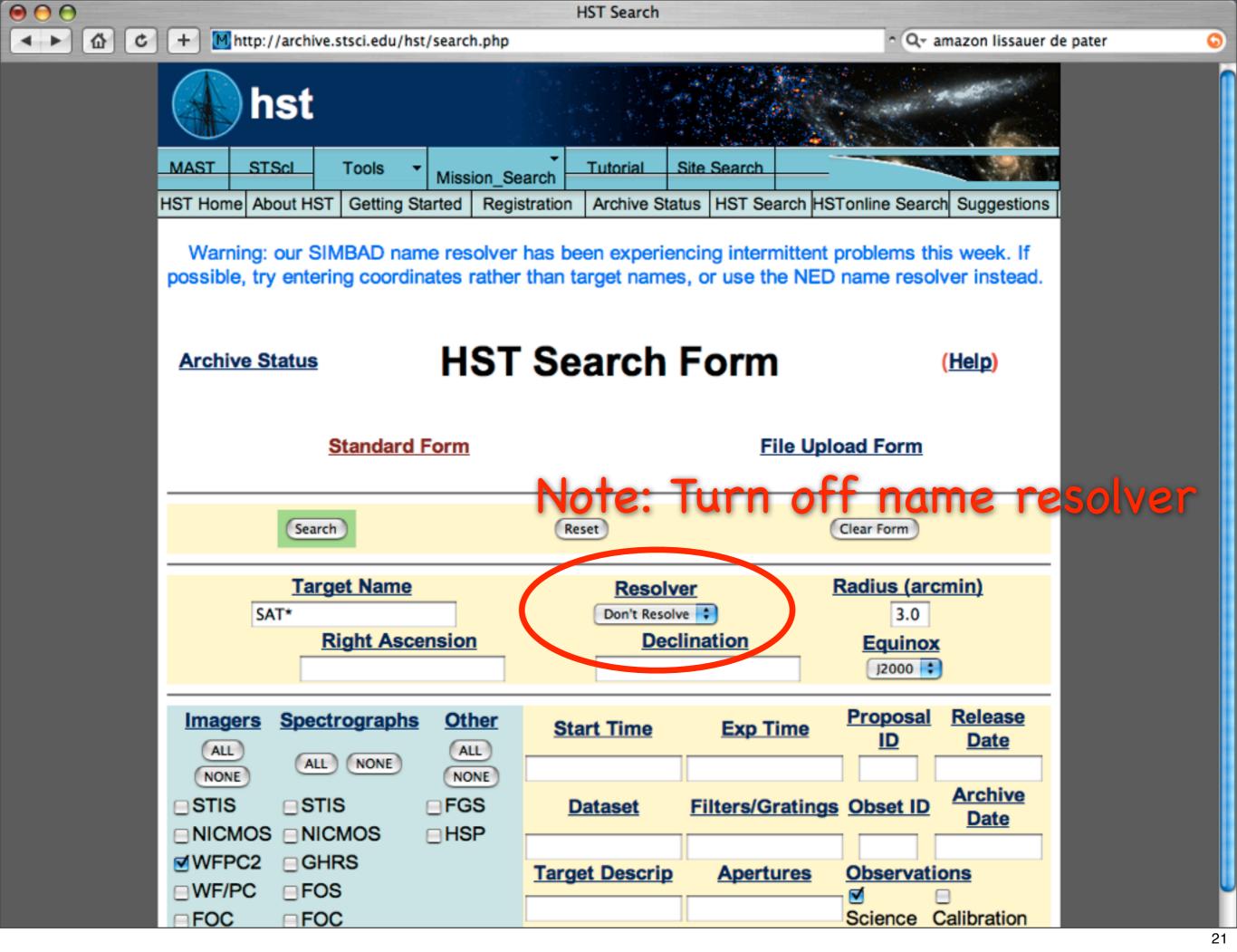
In this example we get the state (STATE) of Titan as seen from the Huygens probe at the UTC epoch 2004 NOV 21 02:40:21.3. The state vector is returned in the J2000 inertial reference frame and has been corrected for both light time and stellar aberration (LT+S). The one-way light time (LT) is also returned.

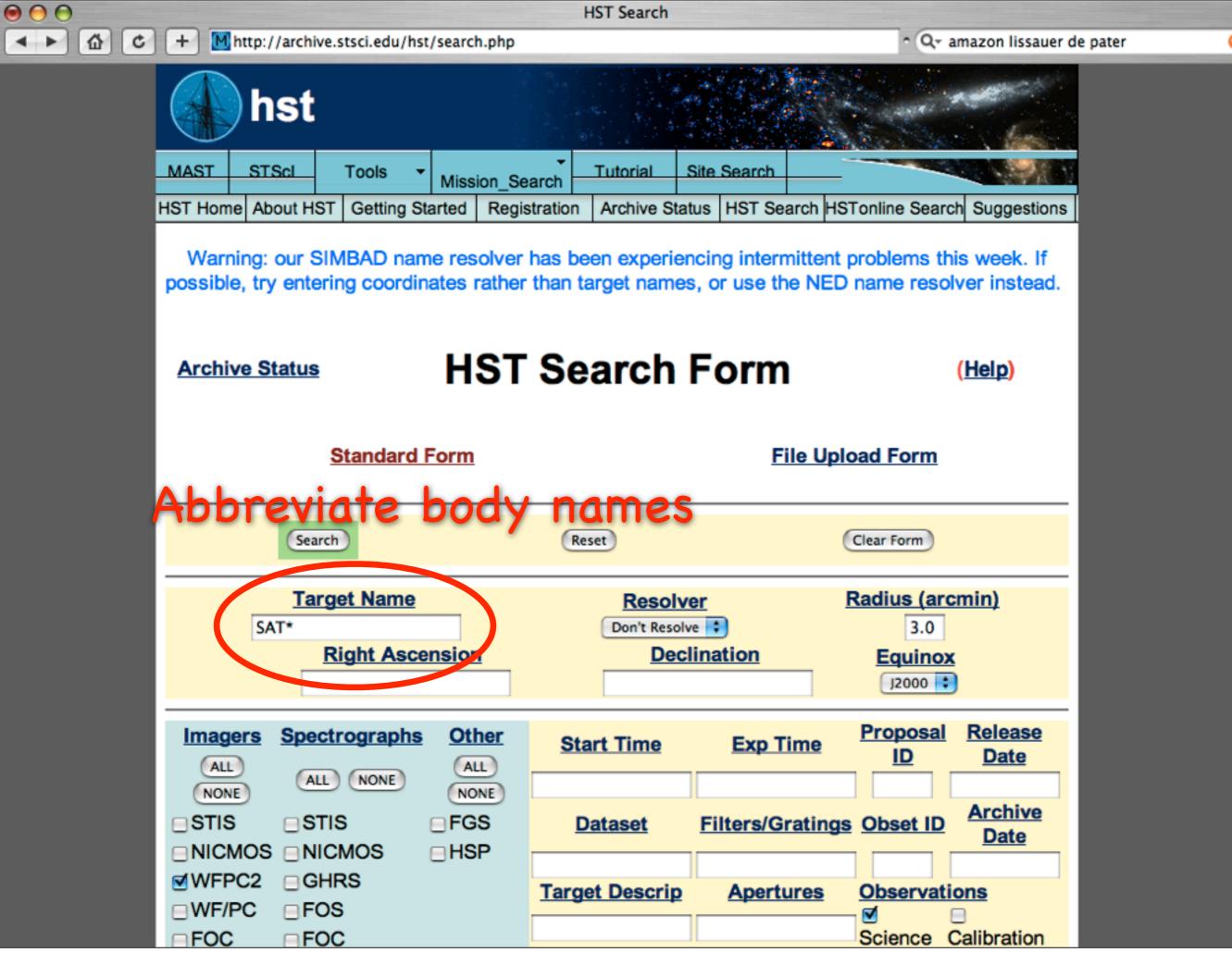
A SPICE leapseconds file (NAIF0008.TLS) is used, as is a SPICE ephemeris file (HUYGENS_3_MERGE.BSP) containing ephemeris data for the Huygens probe (-150), Saturn barycenter (6), Saturn mass center (699), Saturn's satellites (6xx) and the sun (10), relative to the solar system barycenter.

Space Telescope Archive

- Designed primarily for astrophysical use.
 - Not ideal for planetary searches.
 - ...but still usable!
- archive.stsci.edu/hst/search.php

















Mhttp://archive.stsci.edu/hst/search.php







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Plot marked spectra

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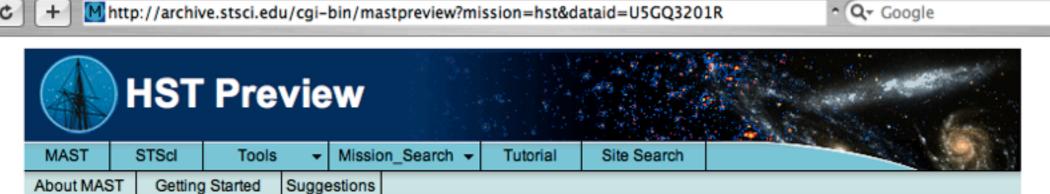
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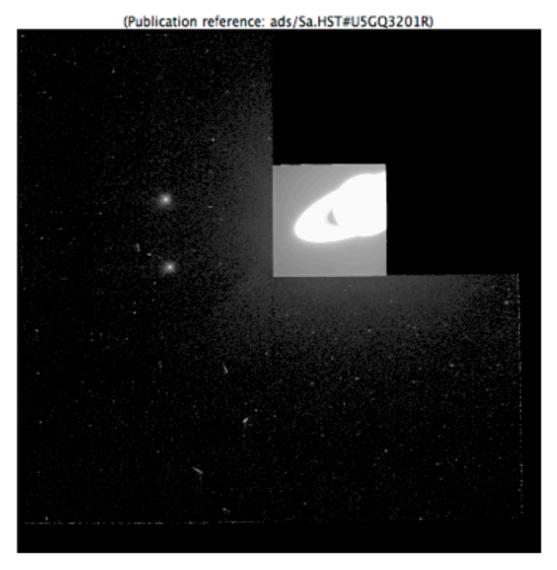
Mark proprietary

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<u>Mark</u>	<u>Dataset</u>	<u>Target Name</u>	RA (J2000)	Dec (J2000)	Ref	Start Time	Stop Time	Exp Time	<u>lr</u>
	U5GQ3201R	SATURN-EAST	02 49 6.86	+13 30 52.1	<u>14</u>	1999-11-03 19:47:00	1999-11-03 19:47:00	30.000	W
	U5GQ3202M	SATURN-EAST	02 49 6.82	+13 30 51.9	<u>14</u>	1999-11-03 19:50:00	1999-11-03 19:50:00	5.000	W
	U5GQ3203R	SATURN-EAST	02 49 6.78	+13 30 51.7	<u>14</u>	1999-11-03 19:53:00	1999-11-03 19:53:00	0.500	W
	U5GQ3204R	SATURN-EAST	02 49 6.74	+13 30 51.6	<u>14</u>	1999-11-03 19:56:00	1999-11-03 19:56:00	0.400	W
	U5GQ3205R	SATURN-EAST	02 49 6.70	+13 30 51.4	<u>14</u>	1999-11-03 19:59:00	1999-11-03 19:59:00	0.400	W
	U5GQ3206R	SATURN-EAST	02 49 6.66	+13 30 51.2	14	1999-11-03 20:02:00	1999-11-03 20:02:00	40.000	W
	U6EM2201M	SATURN-EAST	04 54 11.93	+20 46 58.9	<u>13</u>	2001-09-08 04:49:00	2001-09-08 04:49:00	30.000	W
	<u>U6EM2202M</u>	SATURN-EAST	04 54 11.95	+20 46 58.9	<u>13</u>	2001-09-08 04:52:00	2001-09-08 04:52:00	4.000	W
	<u>U6EM2203M</u>	SATURN-EAST	04 54 11.97	+20 46 58.9	<u>13</u>	2001-09-08 04:55:00	2001-09-08 04:55:00	0.400	W
	<u>U6EM2204M</u>	SATURN-EAST	04 54 11.99	+20 46 59.0	<u>13</u>	2001-09-08 04:58:00	2001-09-08 04:58:00	0.350	W
								Evm	



Preview for U5GQ3201R



Preview calibrations are uncertain so preview data should be used for diagnostic/quick-look purposes only.

FITS format More preview format options

Exposure Information

Target Name: SATURN-EAST RA: 02 49 06.86

Observation Date: Nov 3 1999 7:47PM Exp Time: 30

Instrument: WFPC2 Filter/Grating: F336W

SWIFT

- Free, supported integration tool for N-body simulations.
 - Non-colliding bodies only.
- www.boulder.swri.edu/~hal/swift.html

SWIFT



A solar system integration software package

The SWIFT subroutine package written by <u>Hal Levison</u> and Martin Duncan is designed to integrate a set of mutually gravitationally interacting bodies together with a group of test particles which feel the gravitational influence of the massive bodies but do not affect each other or the massive bodies. Four integration techniques are included:

- Wisdom-Holman Mapping^[1] (WHM). This algorithm was created by Jack Wisdom & Matt Holman (Click here for a reference).
- Regularized Mixed Variable Symplectic (RMVS) method. This handles close approachs between test particles and planets. This algorithm was created by Hal Levison & Martin Duncan (Click here for a reference) and based on WHM.
- A fourth order T+U Symplectic (TU4) method. This algorithm was created by Jeff Candy and W. Rozmus. Also see Martin Duncan, Brett Gladman, and Jeff Candy (Click here for a reference).
- A Bulirsch-Stoer method.

The package is designed so that the calls to each of these look identical so that it is trivial to replace one with another.

We know that SWIFT will run on HP9000s, SUNs, MIPS, IBM RS6000, DEC alphas (UNIX and Linux), PC's (Linux) and DECstations. We do not support VMS, but Sandy Keiser of DTM have written a version for VMS, which we are distributing. Click here for more information of the VMS version. If you are going to run SWIFT on a machine not listed, please let us know how it works.

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Light Scattering Codes

- Mie Scattering (spheres):
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- T-matrix models (irregular particles):
 - www.giss.nasa.gov/~crmim/t_matrix.html
- Discrete dipole approximation (arbitrary particle shapes of limited size):
 - www.astro.princeton.edu/~draine/DDSCAT.6.0.html









http://web.mit.edu/cegeon/Public/Classes/12815/Mie/

~ Q- miev mie scattering



Index of /cegeon/Public/Classes/12815/Mie

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-				
2	Parent Directory	16-Oct-2003 16:19	-	
?	Compile_Run	09-Aug-1999 15:30	1k	
?	<u>ErrPack.f</u>	09-Aug-1999 15:30	3k	
?	ErrPack.o	14-0ct-2003 12:23	6k	
?	MIEV.doc	09-Aug-1999 15:30	21k	
?	MIEV0	15-Oct-2003 00:18	897k	
? ?	MIEV0.f	09-Aug-1999 15:30	63k	
?	MIEV0.o	14-0ct-2003 12:23	97k	
?	<u>Makefile</u>	09-Aug-1999 15:30	1k	
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<u> </u>	ex3-x100n1.txt	14-Oct-2003 17:34	9k	
	ex3-x1n001.txt	14-0ct-2003 17:47	9k	
?	g.f	15-Oct-2003 00:17	2k	
I	g.txt	15-Oct-2003 00:12	14k	
	<u>g100.txt</u>	15-Oct-2003 00:18	14k	
	ouput001.txt	14-Oct-2003 12:35	28k	
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	output01.txt	14-Oct-2003 12:43	28k	
<u> </u>	output1.txt	14-Oct-2003 12:42	28k	
?	phs_fnc.f	14-0ct-2003 23:05	3k	
?	<u>qscat.f</u>	14-Oct-2003 12:44	2k	
?	<u>qscat.o</u>	14-0ct-2003 12:44	6k	
1	x100n01.txt	14-0ct-2003 23:07	9k	



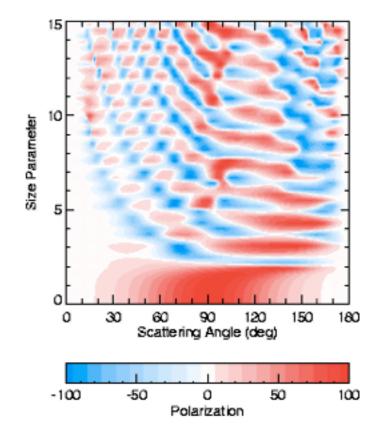
T-Matrix Codes for Computing Electromagnetic Scattering by Nonspherical and Aggregated Particles

By Michael I. Mishchenko, Larry D. Travis, and Daniel W. Mackowski

This web site provides free public access to four *T*-matrix codes for the computation of electromagnetic scattering by homogeneous, rotationally symmetric nonspherical particles in fixed and random orientations, a superposition *T*-matrix code for randomly oriented two-sphere clusters with touching or separated components, and superposition *T*-matrix codes for multi-sphere clusters in fixed and random orientations. All codes are written in Fortran-77. Each code is extensively documented and provides all necessary references to relevant publications.

The double-precision and extended-precision versions of the regular *T*matrix codes are essentially identical. The extended-precision versions are
a factor of 5-8 slower than their double-precision equivalents, but allow
computations for larger (a factor of 2-3) particles. The extended-precision
codes have a more detailed documentation of all the subroutines used.

The regular *T*-matrix codes are applicable to rotationally symmetric particles with equivalent-sphere size parameters exceeding 100. At present, the *T*-matrix method is the fastest exact technique for the computation of nonspherical scattering based on a direct solution of Maxwell's equations. The *T*-matrix codes are orders of magnitude faster than those based on the DDA, VIEF, and FDTD techniques.



The regular *T*-matrix codes for randomly oriented particles are based on the analytical orientation averaging procedure described in the paper M. I. Mishchenko, J. Opt. Soc. Am. A 8, 871-882 (1991). This efficient procedure makes the codes 1 to 2 orders of magnitude faster than *T*-matrix codes based on the standard numerical averaging approach. The superposition bisphere and multi-sphere *T*-matrix codes are based on similar analytical approaches and are described in the papers M. I. Mishchenko and D. W. Mackowski, Opt. Lett., vol. 19, 1604- 1606 (1994) and D. W. Mackowski and M. I. Mishchenko, J. Opt. Soc. Amer. A., vol. 13, 2266-2278 (1996). In application to bispheres, the multi-sphere *T*-matrix code is slower than the bisphere code. However, it can be applied to clusters with a number of components larger than 2.

A general **review** of the *T*-matrix method can be found in the paper M. I. Mishchenko, L. D. Travis, and D. W. Mackowski, J. Quant. Spectrosc. Radiat. Transfer 55, 535-575 (1996) (available in the electronic form at www.giss.nasa.gov/~crmim/publications). To receive a reprint of this paper, leave a message at crmim@giss.nasa.gov indicating your name and mailing address. A detailed **user guide** to the regular random-orientation *T*-matrix codes was also published [JQSRT 60, 309-324 (1998)] and is available here in the pdf format. A **user guide** to the *T*-matrix codes for

[Back to B.T. Draine's home page.]

The Discrete Dipole Approximation for Scattering and Absorption of Light by Irregular Particles

DDSCAT, a Fortran code for calculating scattering and absorption of light by irregular particles, has been jointly developed by Bruce T. Draine (Dept. of Astrophysical Sciences, Princeton University) and Piotr J. Flatau (Scripps Institution of Oceanography, UCSD); the current version is DDSCAT.6.0. This code is publicly available (see below). If you choose to use it, please send email to cdraine@astro.princeton.edu "registering" as a user; registered users of DDSCAT will be notified when updates to the code are made.

User Guide for DDSCAT.6.0

An extensive User Guide is available: "User Guide to the Discrete Dipole Approximation Code DDSCAT.6.0", by B.T. Draine and Piotr J. Flatau. A copy of this document (postscript or pdf) can be obtained from astro-ph/0309069 and can be cited as

Draine, B.T., and Flatau, P.J. 2003, "User Guide to the Discrete Dipole Approximation Code DDSCAT.6.0", http://arxiv.org/abs/astro-ph/0300969.

The following files are of interest:

- <u>astro-ph/0309069.ps.qz</u>: User Guide for DDSCAT.6.0 (gzipped postscript)
- astro-ph/0309069.pdf: User Guide for DDSCAT.6.0 (pdf)

Downloading the DDSCAT Code

Unix users will find it most convenient to download

<u>ddscat.6.0.tgz</u>: gzipped tarfile containing complete source code and documentation for DDSCAT.6.0

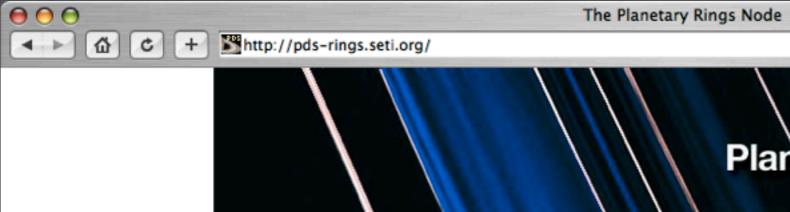
Non-unix users can use anonymous ftp to download the code (*.FOR files), documentation (UserGuide.ps), and sample parameter files (ddscat.par and diel.tab).

History of recent releases:

- DDSCAT.5a9 was released 1998 December 23. It fixes a bug in DDSCAT.5a8, which resulted in incorrect evaluation of elements of the Mueller scattering matrix (other than S_11, which was correct) for scattering planes other than phi=0.
 - A new version of orient.f was released 1999 March 16. Orientational averages are now evaluated as described in the UserGuide; prior to this
 date averaging over cos(Theta) was evaluated by dividing the range of cos(Theta) into NTHETA equal intervals, evaluating the scattering at the
 midpoint of each interval, and taking the mean. With the new version of orient.f, Simpson's rule is now used for the quadrature when an odd
 value of NTHETA is specified by the user.
- DDSCAT.5a10 was released 2000 August 9. It provides a new target option -- NSPHER -- to create targets consisting of the unions of N spheres
 (possibly overlapping) of arbitrary sizes and locations. It also uses a more recent version of the LAPACK code used by subroutine PRINAXIS.
 - A new target option -- PRISM3 -- to generate a triangular prism was added to DDSCAT.5a10 on 2002.02.12

PDS Rings Node pds-rings.seti.org

- Google (TM) search of our site.
- On-line ephemeris and planetary diagramming tools: /tools
- Browsable views of all Cassini and Voyager images: /previews
- Software toolkits in C and FORTRAN: /toolkits
- Voyager and Galileo catalogs, functional but being phased out: /catalog
- New general catalog in development: /demo



Planetary Rings Node

Rings Node Home

Cassini News

Cassini Mission (IPL) Press Release Images Data and Information

Ringed Planets

Jupiter Saturn Uranus Neptune

Missions and Data

Data Search . . .

New Horizons

Cassini

Voyager

Galileo

Hubble Telescope

Saturn RPX 1995

Uranus RPX 2007

Occultations

Astrometry

Downloads

Data volumes Zip archives Previews/Thumbnails

Resources

Tools Toolkits Glossary

Contacts

Mark Showalter Mitch Gordon Neil Heather

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The Rings Node of NASA's Planetary Data System is devoted to archiving, cataloging, and distributing scientific data sets relevant to planetary ring systems.



Google Search

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Recent Highlights

- What's New on line.
- ROSES Support
 - A set of ROSES 2007 support pages providing information and data access to support the rings related analysis programs in this year's announcement is now available.
- Cassini Data
 - Our Cassini Data Archive has been updated for the July, 2007 data releases. New data arrives in October.
 - New: A simplified version of the Cassini CIRS (thermal infrared) data set is in peer review. Comments welcome!
 - Cassini Press Release Images are on line and updated regularly.
 - · Our ephemeris tools are always up to date with the latest Cassini trajectory information.
- Voyager Data
 - New: A complete set of <u>SPICE C kernels</u> is now available for the Voyager 1 Saturn encounter. More kernels will follow shortly.
 - New: Complete, expanded volumes of IRIS (thermal spectra) data are on line and ready for peer review.
 - New: Preliminary versions of the calibrated Uranus images are on line in volumes VGISS_0001, VGISS_0002, and VGISS_0003. Check back for updates.
 - New: The complete set of Voyager images from Saturn in calibrated and geometrically corrected formats is through peer review and now in lien resolution. Comments are still welcome.
 - Our Voyager information pages have received a complete makeover, including the addition of more SPICE data and information.
 - Ring occultation data sets from the PPS (photopolarimeter) and UVS (ultraviolet spectrometer) are complete.
 - The Radio Science (RSS) ring occultation data set is in lien resolution following peer review. Check back for updates.
- Other Data

6

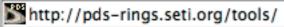














Planetary Rings Node

Rings Node Home

Cassini News

Cassini Mission (JPL) Press Release Images Data and Information

Ringed Planets

Jupiter

Saturn

Uranus

Neptune

Missions and Data

Data Search . . .

New Horizons

Cassini

Voyager

Galileo

Hubble Telescope

Saturn RPX 1995

Uranus RPX 2007

Occultations

Astrometry

Downloads

Data volumes
Zip archives

Previews/Thumbnails

Resources

Tools Toolkits

Glossary

Contacts

Mark Showalter Mitch Gordon

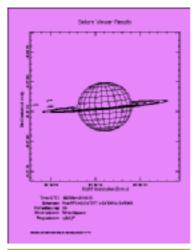
Neil Heather

Submit a comment

Rings Node On-line Tools

A set of tools to assist planetary scientists in the planning, acquisition and interpretation of observations of the ringed (and possibly-ringed) planets.

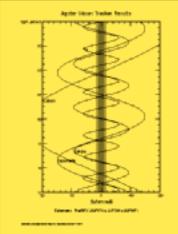
Versions of these tools are now available to support the <u>Cassini Mission</u> and the <u>New Horizons</u> Jupiter flyby.



Planet Viewers

These forms enable you to generate a diagram showing the appearance of a planetary system at a specified time. Bodies and rings are rendered with terminators and shadows as appropriate. The viewpoint can be Earth's center, a particular Earth-based observatory, or a spacecraft.

Click for Jupiter, Saturn, Uranus, Neptune, or Mars.



Moon Trackers

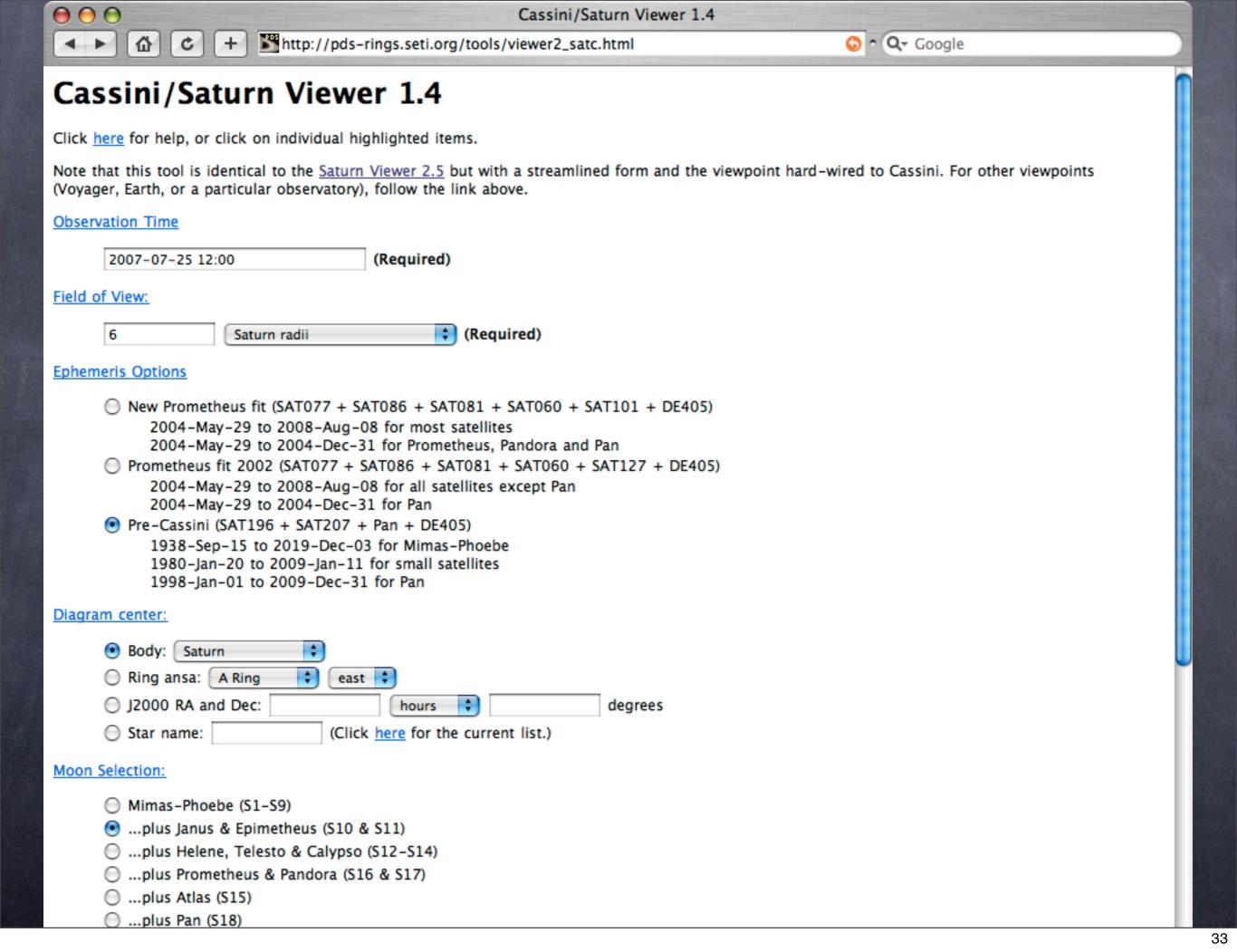
These forms enable you to generate a diagram showing the apparent eastwest motion of one or moons relative to the disk of a planet, within a specified time period.

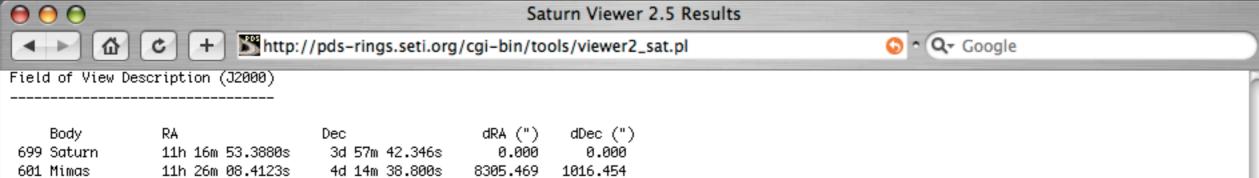
Click for Jupiter, Saturn, Uranus, Neptune, or Mars.

Ephemeris Generators

These forms enable you to generate a table listing useful information about the viewing geometry for a planet and/or any of its moons as a function of time. You are free to specify which of a variety of useful quantities to tabulate (e.g. RA and dec, phase angle, ring opening angle, distance, lunar phase, etc.).

Click for Jupiter, Saturn, Uranus, Neptune, or Mars.





11h 33m 44.8546s 4d 19m 37.833s 15135.742 1315.487 602 Enceladus 603 Tethys 11h 18m 25.2529s 4d 04m 50.235s 1374.680 427.889 604 Dione 11h 03m 09.6365s 3d 38m 09.124s -12326.747 -1173.222 605 Rhea 11h 42m 41.5239s 4d 22m 41.401s 23166.546 1499.055 606 Titan 12h 17m 10.5303s 5d 07m 37.133s 54127.480 4194.787 10h 08m 40.2850s 1d 32m 40.526s -61249.831 -8701.820 607 Hyperion 10d 24m 56.422s -29043.712 23234.076 608 Iapetus 10h 44m 32.5026s 609 Phoebe 12h 23m 40.7994s 3d 36m 01.281s 59967.528 -1301.065 610 Janus 11h 25m 55.9615s 4d 09m 47.664s 8119.153 725.318 11h 01m 40.4823s 3d 37m 03.642s -13660.863 -1238.704 611 Epimetheus

Sub-solar latitude (deg): -11.42301 (-11.45191 to -11.39412)

Ring opening angle (deg): 0.08539 (unlit)

Phase angle (deg): 21.18831

Sub-solar longitude (deg): 200.90068 from ring plane ascending node

Sub-observer longitude (deg): 218.81475

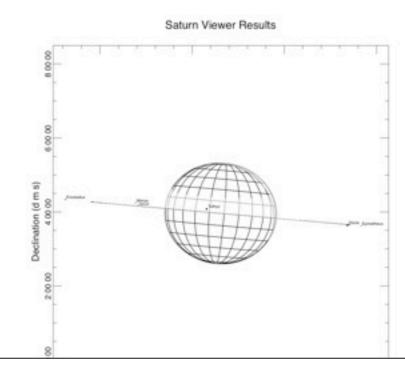
Sun-planet distance (km): 1380.150526 \times 10^6 Observer-planet distance (km): 2.285400 \times 10^6

Light travel time (sec): 7.623273

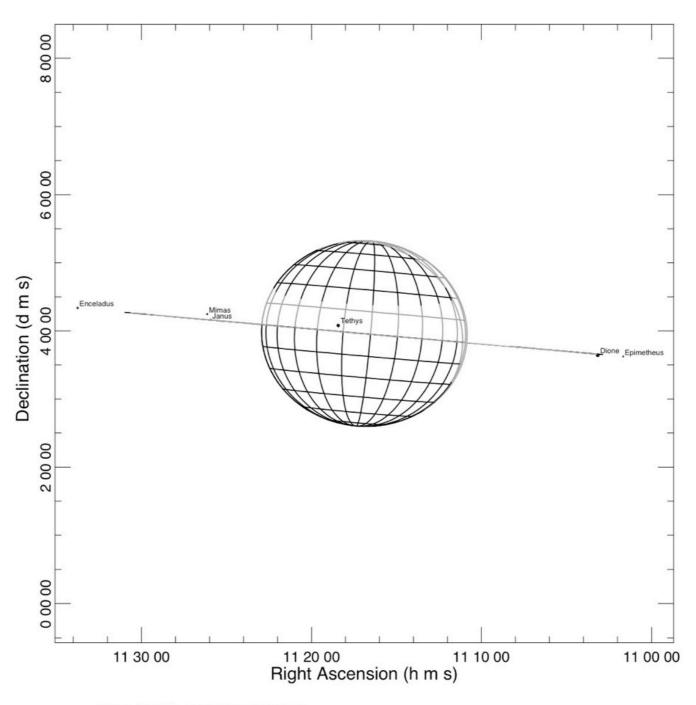
F Ring pericenter (deg): 281.77575 from ring plane ascending node

F Ring ascending node (deg): 152.94911

Preview:



Saturn Viewer Results



Time (UTC): 2007-07-25 12:00

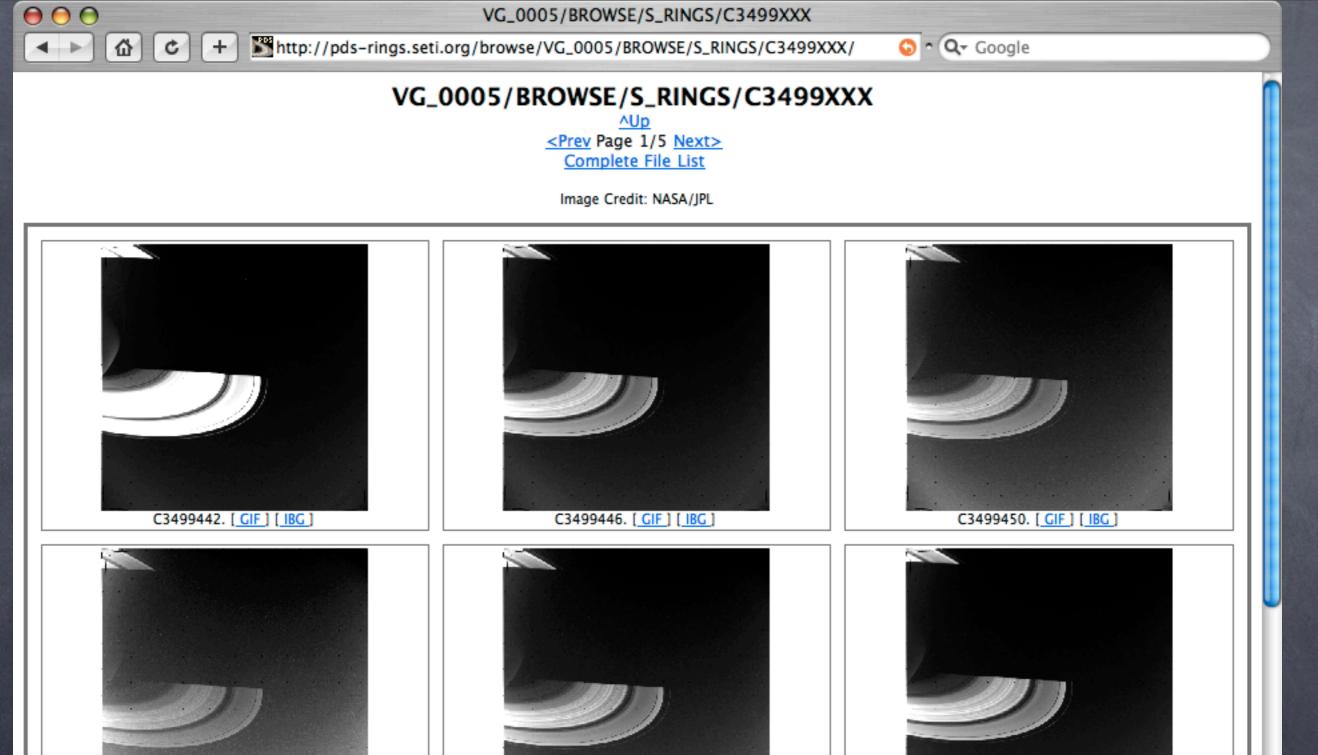
Ephemeris: Pre-Cassini (SAT196 + SAT207 + Pan + DE405)

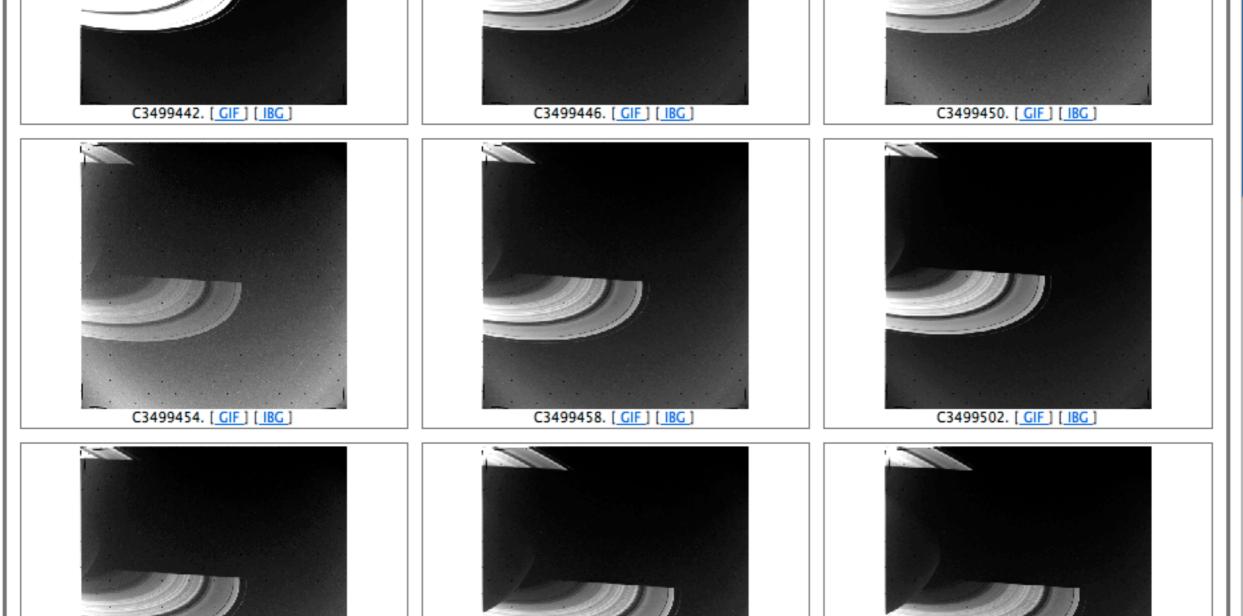
Viewpoint: Cassini

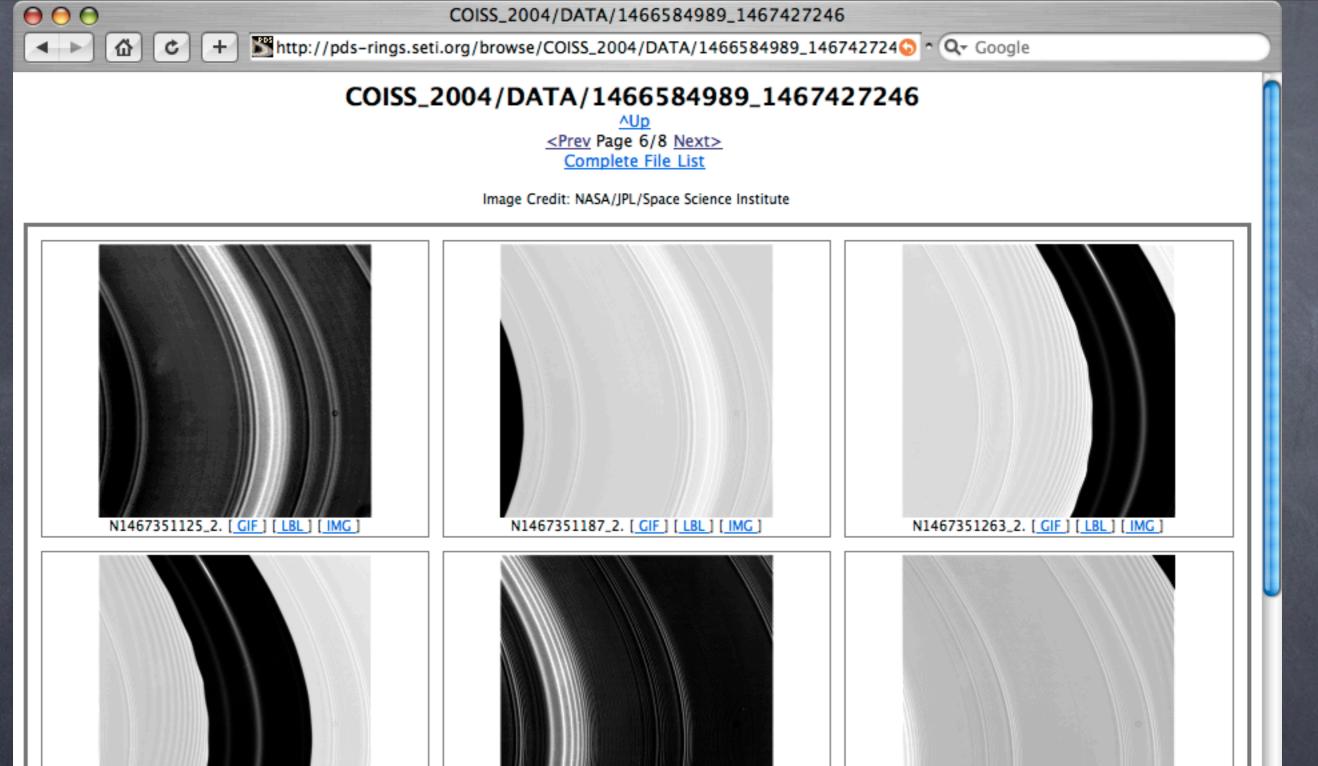
Moon selection: Mimas-Phoebe, Janus & Epimetheus

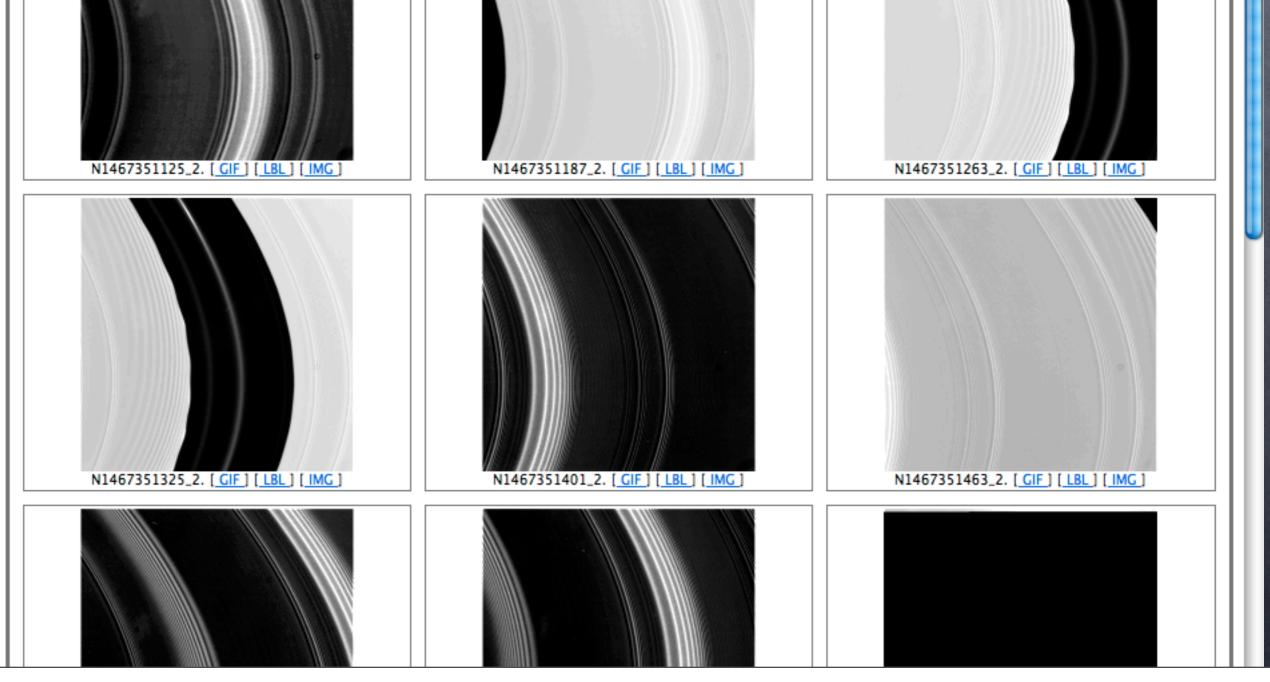
Ring selection: A,B,C,F

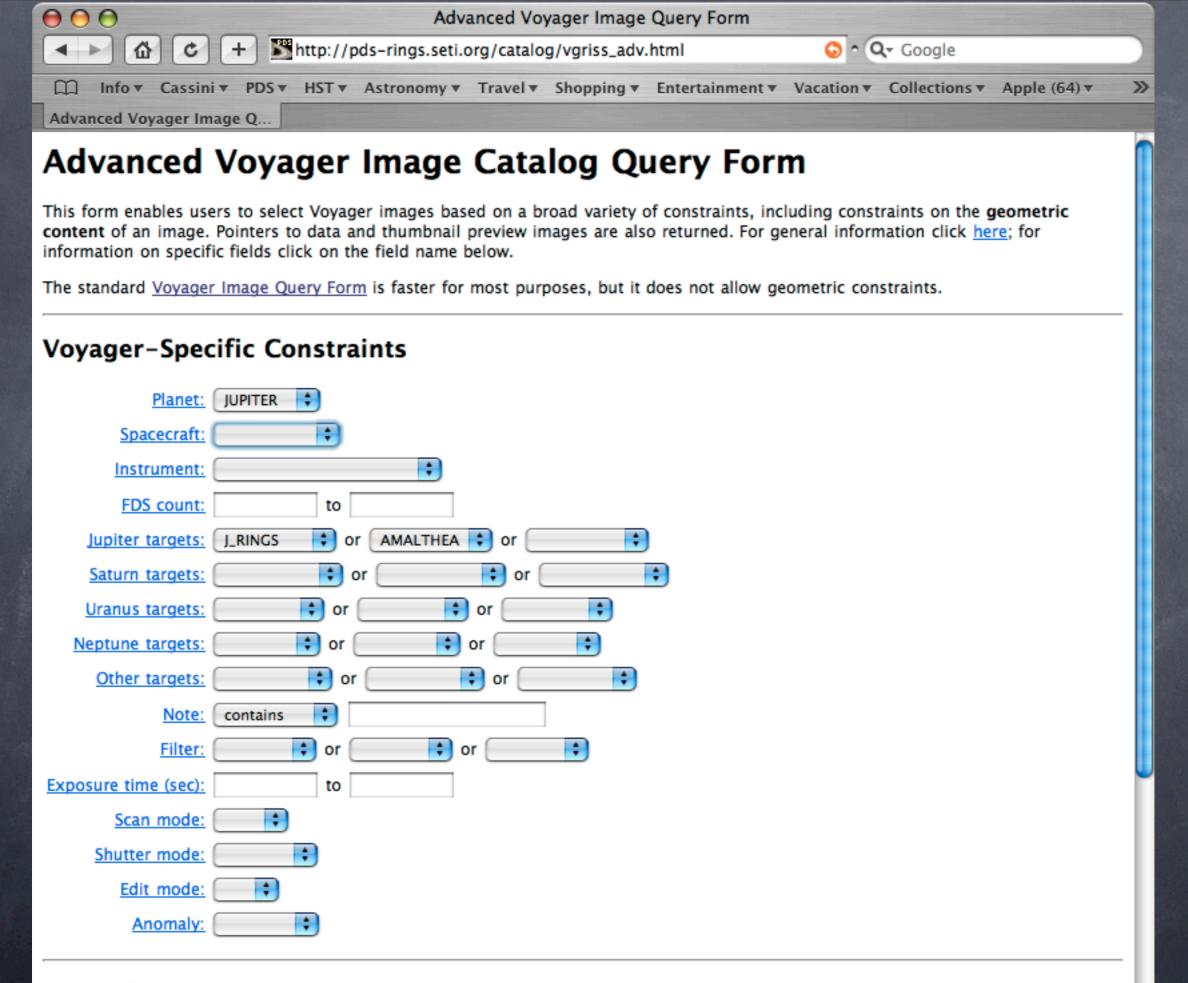
Generated by the Saturn Viewer Tool, PDS Rings Node, Sun Jul 22 07:44:58 2007



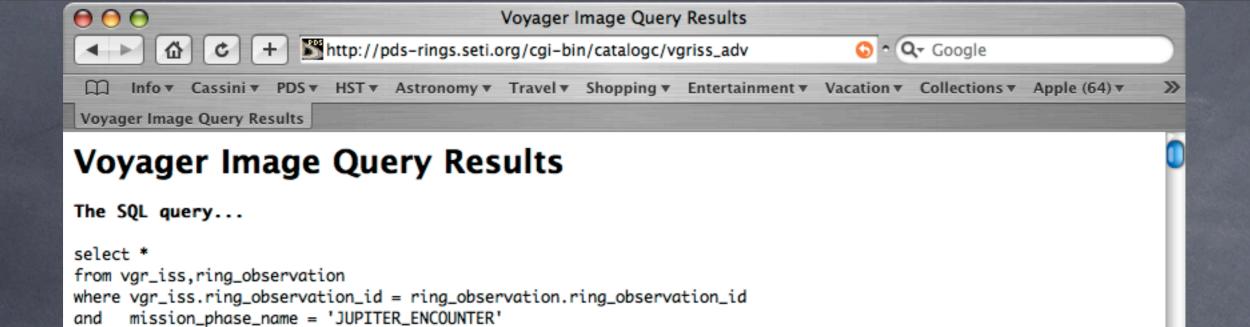








General Constraints



Listing type: Detailed

and target_name in ('J_RINGS', 'AMALTHEA')

order by vgr_iss.ring_observation_id, vgr_iss.occurrence_number

Include previews: No Records to skip: 0

Records found: 132 Records returned: 100

Note: This query listing is limited to 100 matches. You may view the remainder of the list by using the "Skip the first ___ records" option.

Top | Summary listing | Detailed listing | Help

SUMMARY LISTING

Click on record number to jump to detailed entry. Click on FDS count for thumbnail preview image.

#	FDS	Planet	Vgr	Camera	Target	Filter	Texp	Scan	Shutter	Ring?
001	15508.28	JUPITER	1	NARROW	AMALTHEA	CLEAR	0.960	1:1	BSIMAN	NO NO
002	15508.29	JUPITER	1	WIDE	AMALTHEA	CLEAR	0.360	1:1	BSIMAN	NO
003	15933.37	JUPITER	1	NARROW	AMALTHEA	CLEAR	0.960	1:1	NAONLY	NO
004	16079.23	JUPITER	1	NARROW	AMALTHEA	ORANGE	0.960	1:1	BSIMAN	NO
005	16079.25	JUPITER	1	NARROW	AMALTHEA	GREEN	0.720	1:1	BSIMAN	NO
006	16079.27	JUPITER	1	NARROW	AMALTHEA	VIOLET	0.480	1:1	BSIMAN	NO
007	16103.57	JUPITER	1	WIDE	AMALTHEA	CLEAR	0.360	1:1	WAONLY	NO
008	16104.02	JUPITER	1	NARROW	AMALTHEA	CLEAR	0.960	1:1	BSIMAN	NO
009	16104.03	JUPITER	1	WIDE	AMALTHEA	CLEAR	0.360	1:1	BSIMAN	NO
010	16137.11	JUPITER	1	WIDE	AMALTHEA	CLEAR	0.360	1:1	WAONLY	NO
011	16137.16	JUPITER	1	NARROW	AMALTHEA	CLEAR	0.960	1:1	BSIMAN	NO

